LIGO

# Advanced LIGO
# The Next Generation

**Philip Lindquist, Caltech**

**XXXVIII Moriond Conference**

**Gravitational Waves and Experimental Gravity**

March 23, 2003

# The LIGO Mission

## To develop the field

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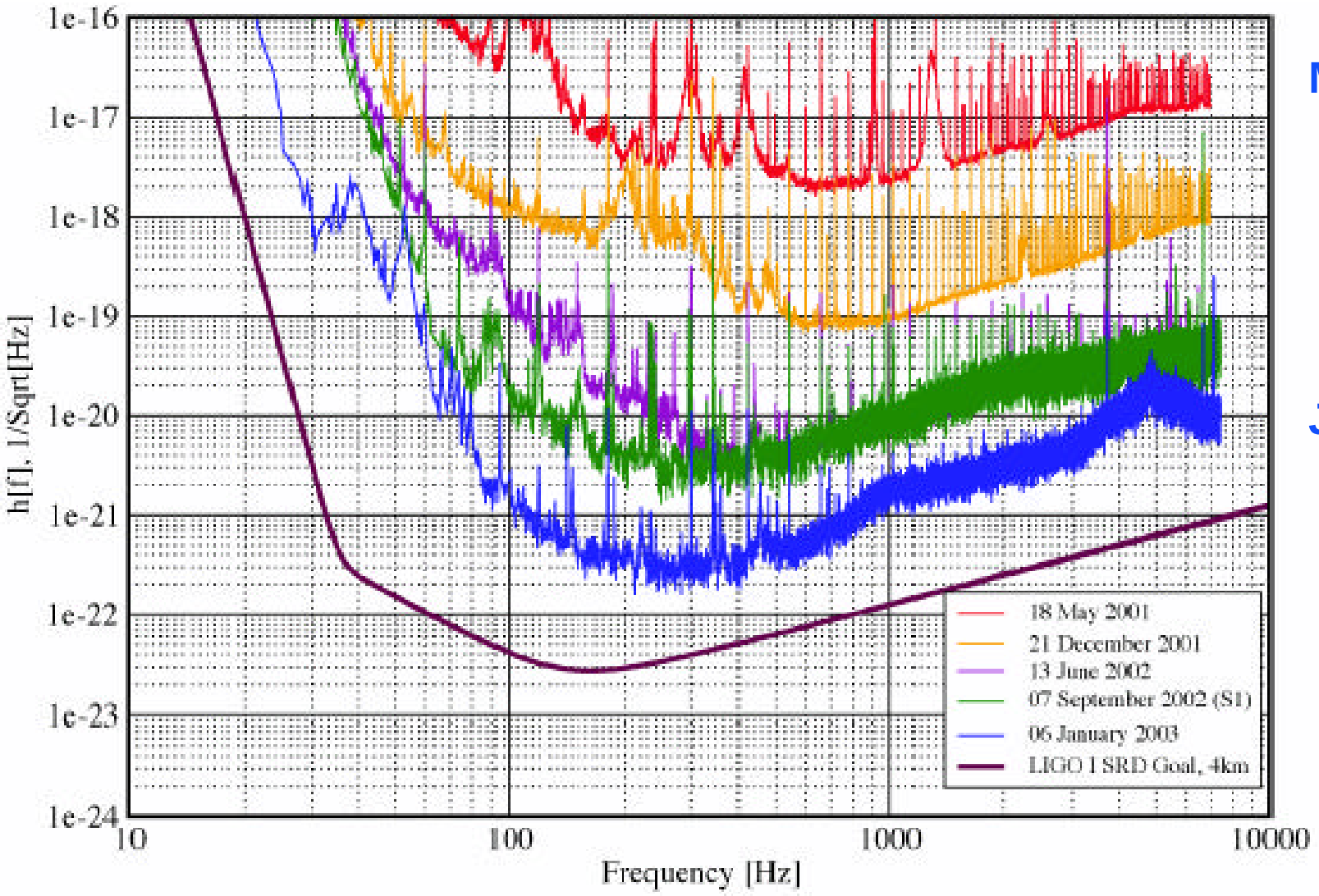
# Livingston 4km Sensitivity History




**May 01**

**Jan 03**

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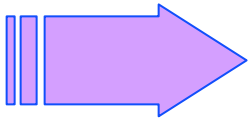
- Detection is plausible with the initial LIGO
- With or without detection, astrophysical community will want/demand more sensitive detectors



**Advanced LIGO**

# Advanced LIGO Reach

- Next Detector
  - » Must be relevant for astrophysics
  - » Should approach limits of reasonable extrapolations for detector physics and technologies
  - » Should lead to realizable, practical, reliable instrument
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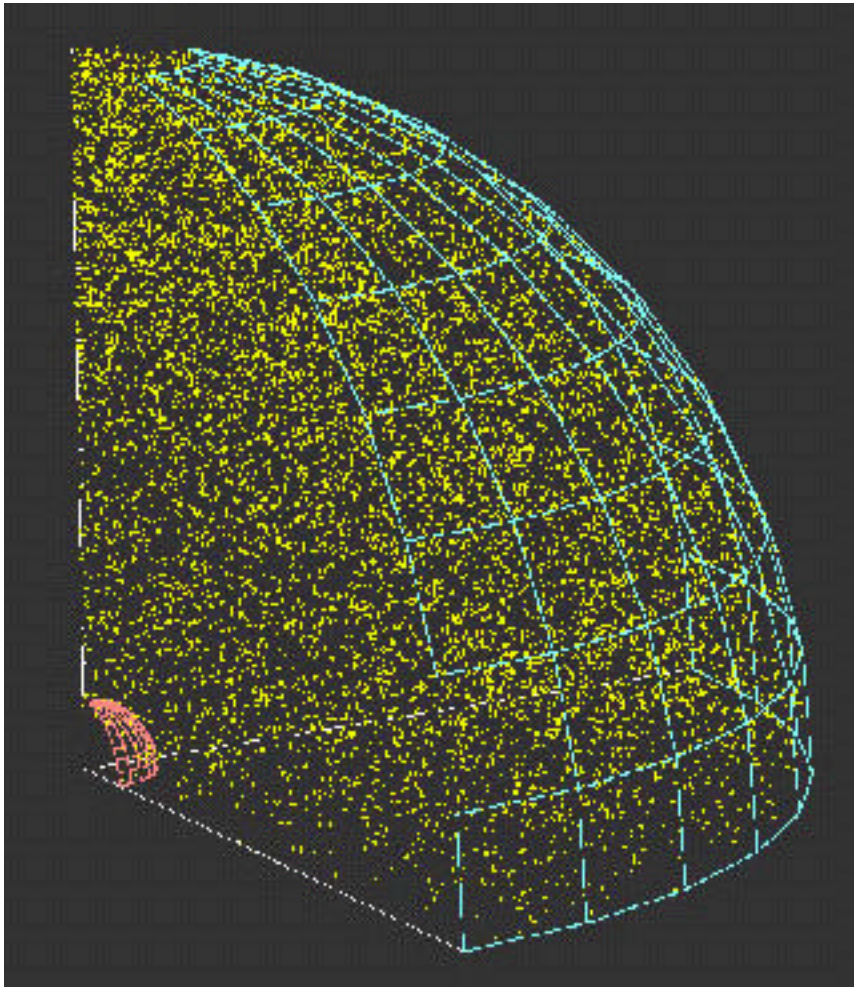
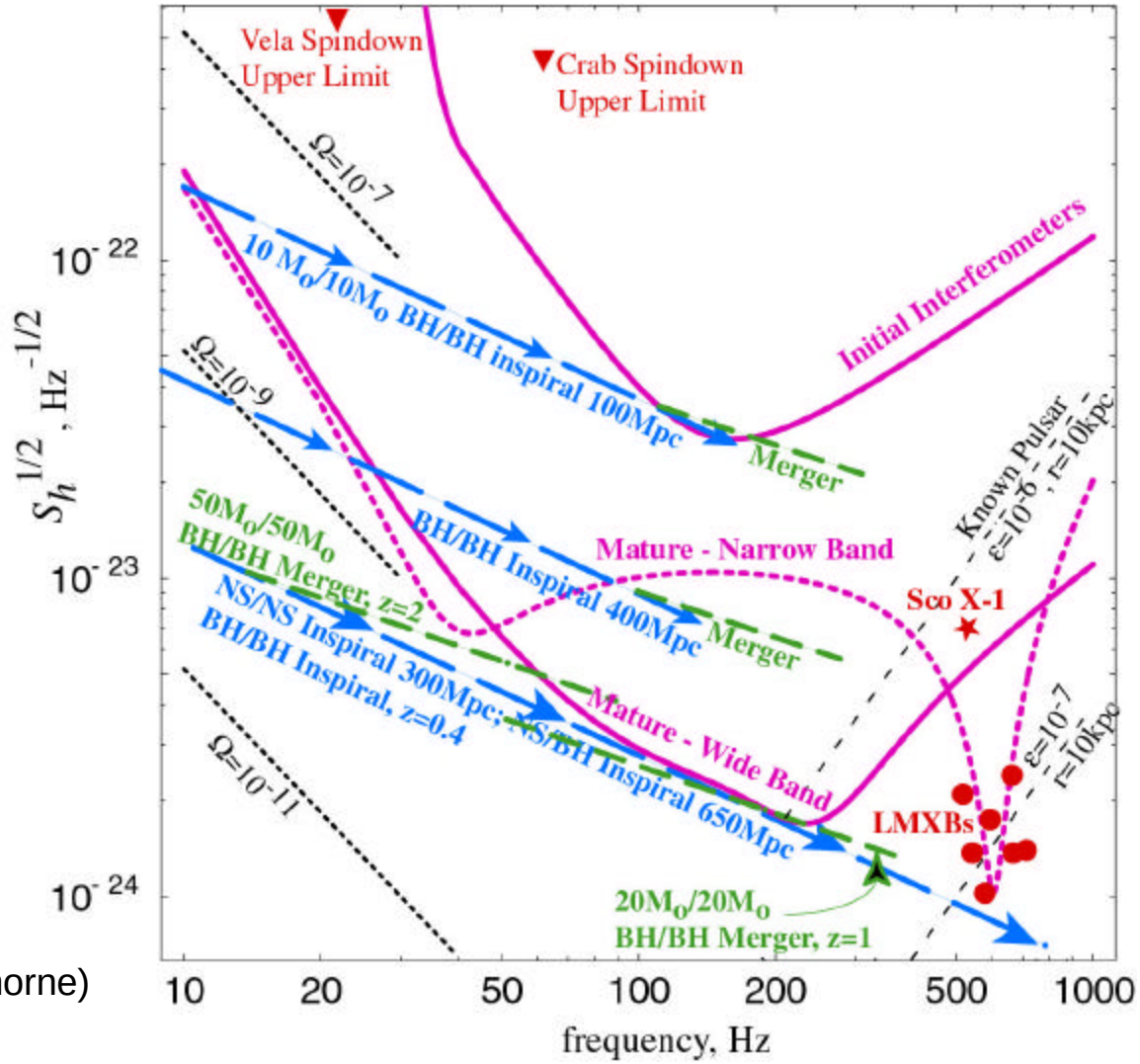
- Advanced LIGO
  - » >10 X sensitivity, ~ 3000 in rate (population density dependent)
  - » **~2.5 hours = 1 year of initial LIGO**



# Astrophysical Reach

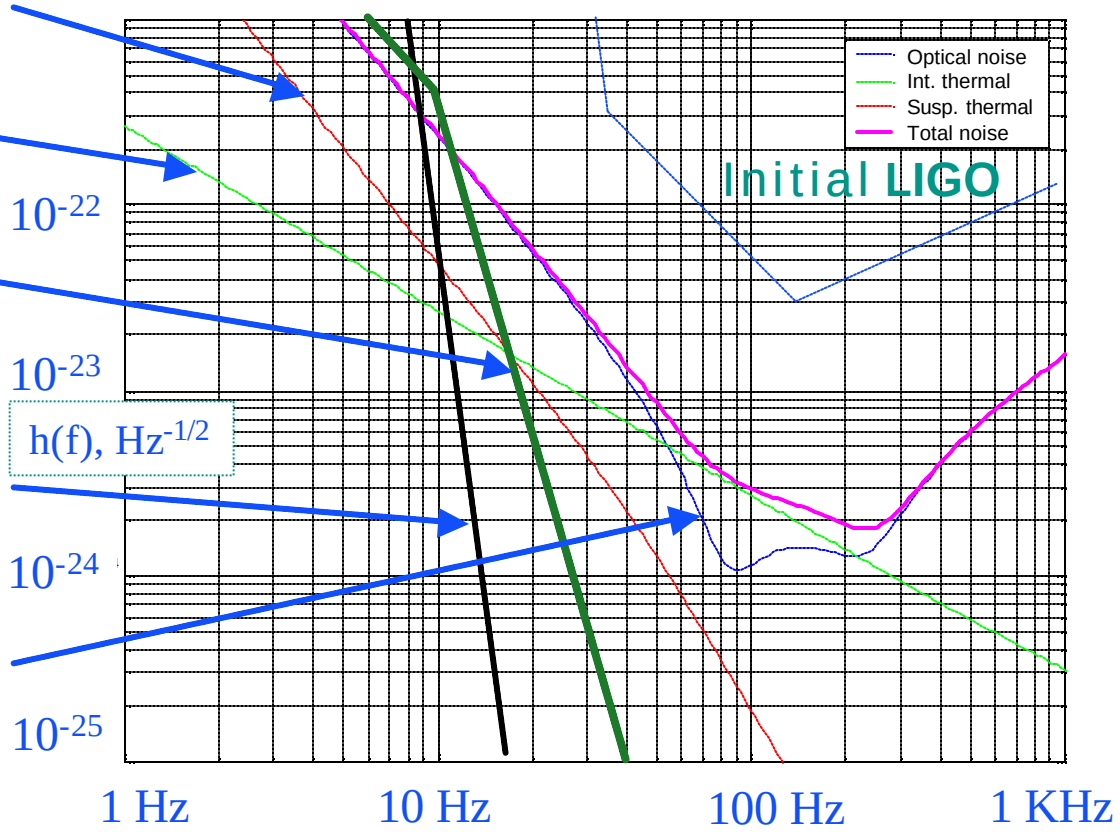
- Neutron Star and Black Hole Binaries
- Spinning Neutron Stars
- NS Birth (Super Novae, AIC)
- Stochastic Background




(Chart by Kip Thorne)

# Projected Advanced LIGO Performance

- **Suspension thermal noise**
- **Internal thermal noise**
- **Newtonian background, estimate for LIGO sites**
- **Seismic "cutoff" at 10 Hz**
- **Unified quantum noise dominates at most frequencies for full-power, broadband tuning**

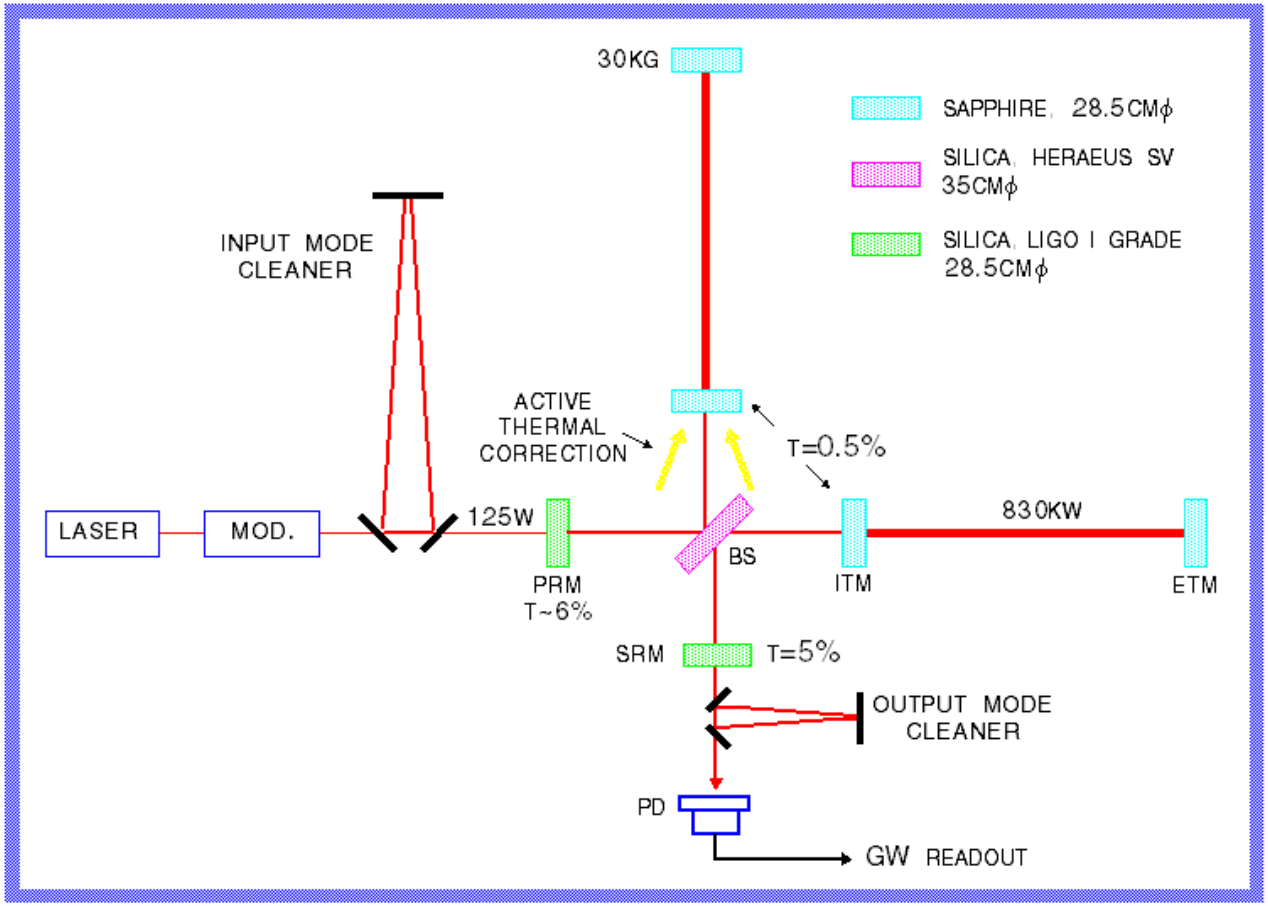


# Advanced LIGO Basic Configuration

**Current LIGO**

**Power recycled**

**Fabry-Perot to increase storage time**



# Advanced LIGO Basic Configuration

**Current LIGO**

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**Advanced LIGO**

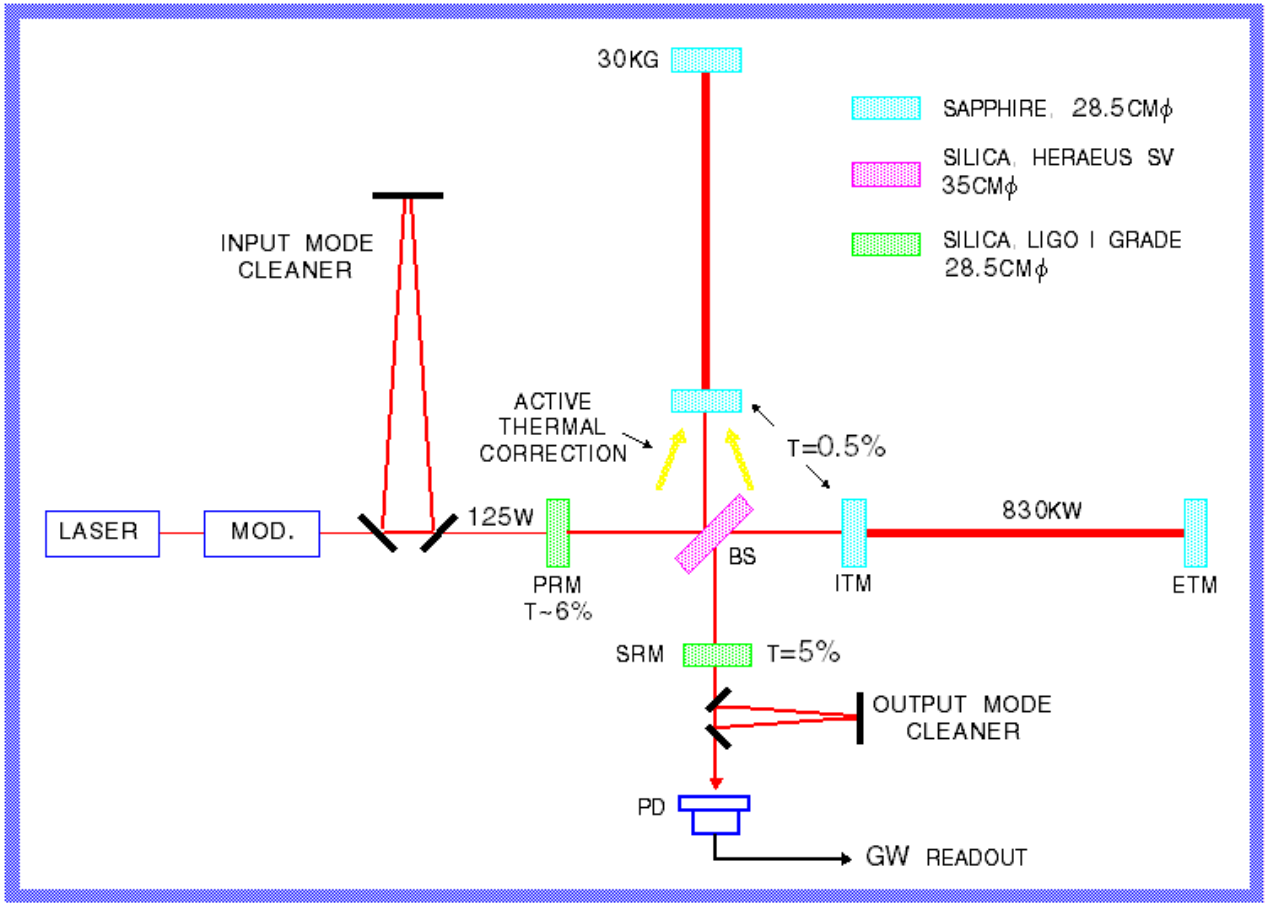
**Increased Power**

**Heavier Test Masses**

**Quad Suspensions**

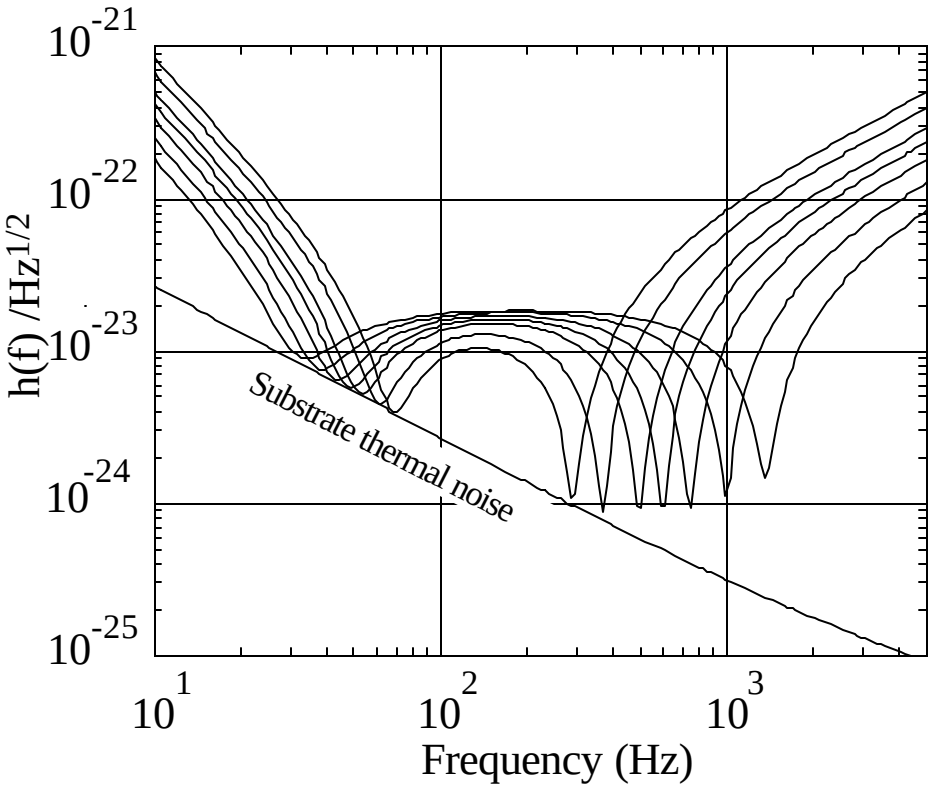
**Improved Isolation**

**Signal Recycling**



# Signal Recycling

- Can focus sensitivity where needed
  - » Sub-wavelength adjustments of resonance in signal recycling cavity
- Allows optimization
  - » Technical constraints
  - » Astrophysical signatures

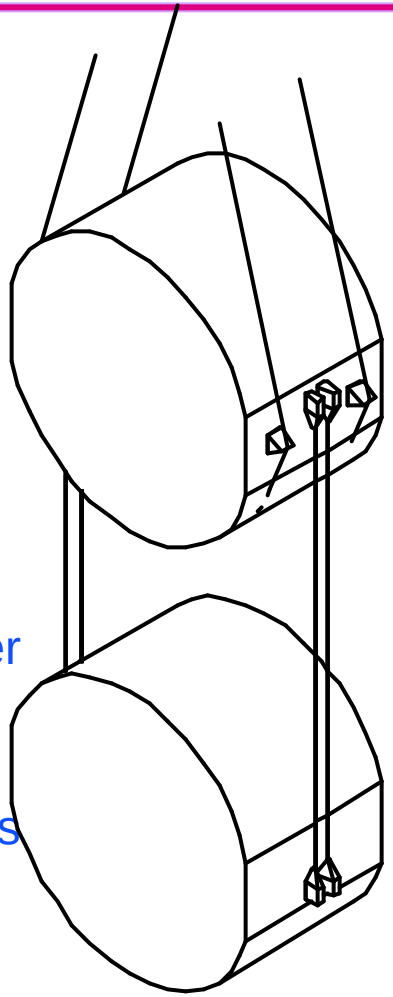


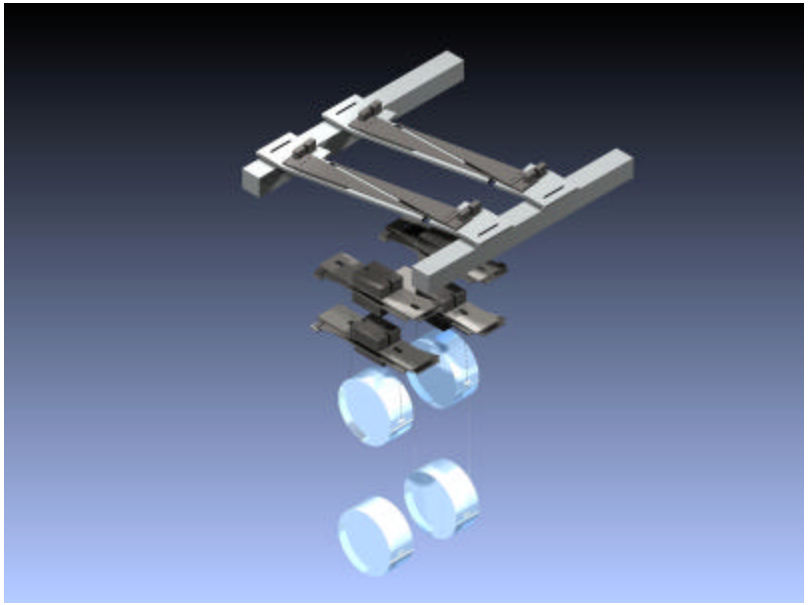
# Limits to Sensitivity--Thermal Noise

- Thermal motion is proportional to $\sqrt{L_{mechanical}}$

Low-loss materials and techniques needed

  - » Test masses: crystalline sapphire, 40 kg, 32 cm diameter
  - » Suspensions: fused silica
  - » Monolithic final stages
  - » Multiple pendulums for control and seismic attenuation **(GEO 600)**

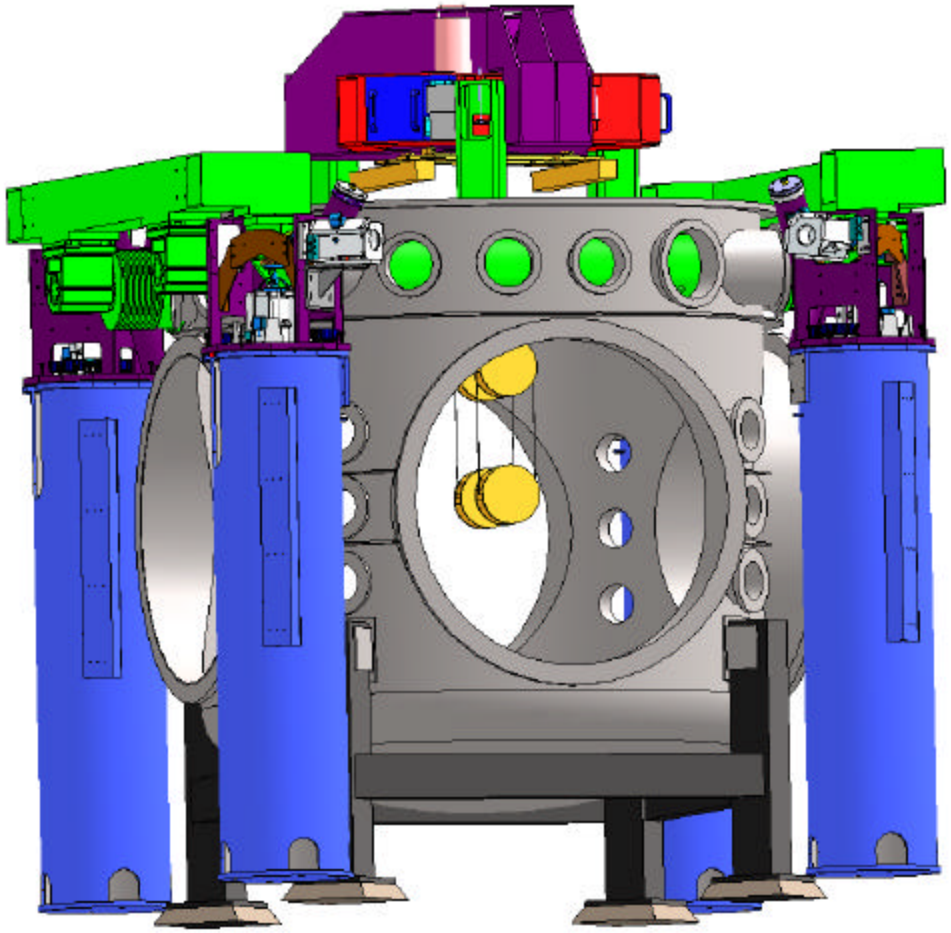




**Optical coating also source of mechanical loss, development underway**

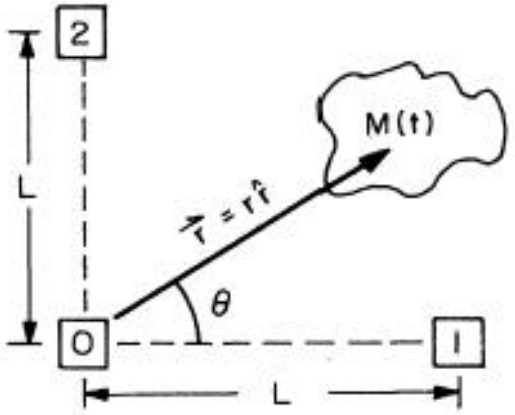
# Limits to Sensitivity--External Forces

- **Coupling of seismic noise through isolation system suppressed using active servo controls, passive "pendulum" isolation**
  - » Two 6-degree-of-freedom platforms stabilized from 0.03 to 30 Hz
  - » Net suppression of motion in gravitational-wave band is 13 orders of magnitude or more
  - » Suppression of motion below the band also critical to hold sensing system (control) in linear domain, avoid up-conversion



# Low-Frequency Limit

- Newtonian background is limit for ground-based detectors (~10 Hz)
  - » Time-varying distribution of mass in vicinity of test mass
  - » Seismic compression, rarefaction of earth dominates
  - » Advanced LIGO targeted to reach this limit for our sites
- For GW astrophysics below 10 Hz, space-based instruments needed ® **LISA**



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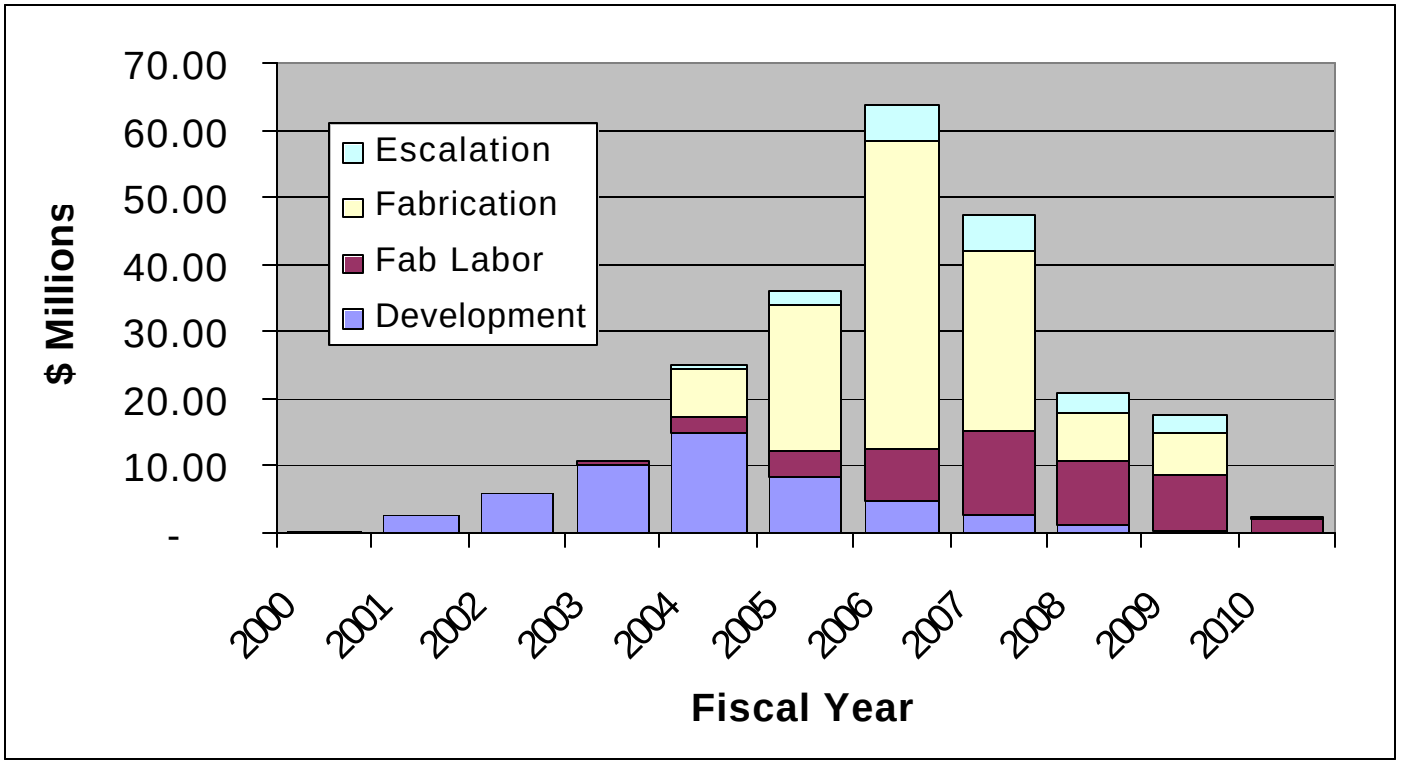
- **Total Cost Estimate: $240,000,000**
- **Proposed LIGO Cost $122,000,000 (fabrication plus some salaries for installation, specialty engineering, added outreach)**
- **$25.5 million provided by collaborators**
  - » **GEO (Hanover, Birmingham, Rutherford, Glasgow)**
  - » **ACIGA**

# Advanced LIGO Timeline

- Initial LIGO Observation 2002 – 2006
  - » 1+ year observation within LIGO Observatory
  - » Significant networked observation with GEO, LIGO, TAMA
- Structured R&D program to develop technologies 1998 - 2005
  - » Conceptual design developed by LSC in 1998
  - » Current Cooperative Agreement carries R&D to Final Design, 2005
- Proposal submitted in Feb 2003 for fabrication, installation
- Long-lead purchases planned for 2004
  - » Sapphire Test Mass material, seismic isolation fabrication
  - » Prepare a ‘stock’ of equipment for minimum downtime, rapid installation
- Start installation in 2007
  - » Baseline is a staged installation, Livingston followed by Hanford Observatories
- Start coincident observations in 2009

# Proposed Funding Profile

- Long lead procurements begin in 2004
- Procurements peak in 2006
- Installation begins in 2007



# The Advanced LIGO Community

- Scientific impetus, expertise, and development provided by LIGO Scientific Collaboration (LSC)
  - » Synergy, critical mass (400+ individuals, 100+ graduate students, 40+ institutions)
  - » International support and significant material participation
  - » Especially strong collaboration with German-UK GEO group, capital partnership
- Advanced LIGO design, R&D, and fabrication shared with participants
  - » LIGO laboratory leads, coordinates, is responsible for observatories
- Continuing support from NSF at all levels
- International network growing: VIRGO, GEO-600, TAMA, ACIGA

# Advanced LIGO Status Summary

- **Initial LIGO is in operation**
  - » We are preparing publications from first science run
  - » Second science run underway
  - » Third science run at target sensitivity scheduled to begin next year
  - » Discovery plausible
- **Advanced LIGO on the horizon**
  - » Advanced R&D and baseline design proceeding
  - » Strong international partnership—GEO, ACIGA
  - » Plan supports start of installation in 2007

**Gravitational Waves:**
**new tool for understanding the Universe,**
**complementary to other observational methods,**
**becoming a reality**

# Acknowledgements


- National Science Foundation
  - NSF PHY-9210038 (Construction and Initial Operations)
  - NSF PHY-0107417 (Current Operations)
- Gary Sanders, LIGO Deputy Director
- David Shoemaker, Advanced LIGO Project Lead