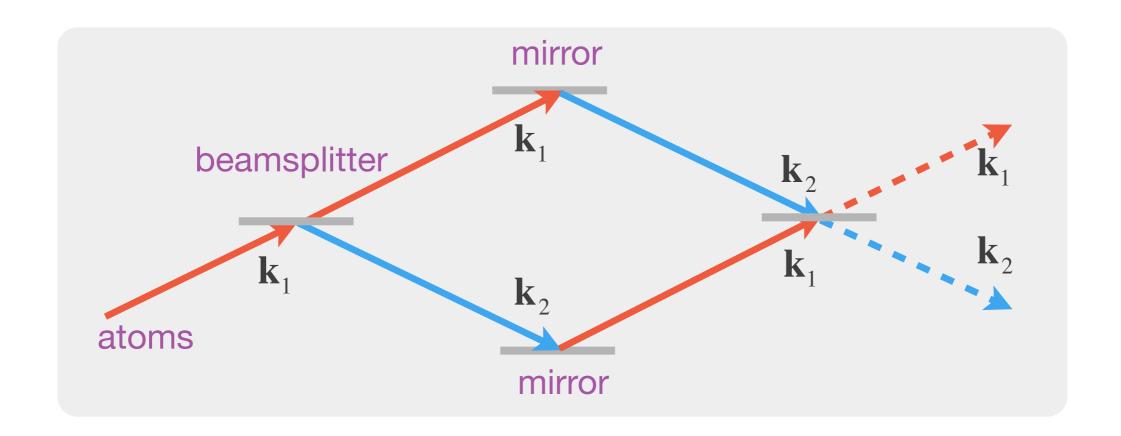
Atom Interferometers

Yanbei Chen California Institute of Technology

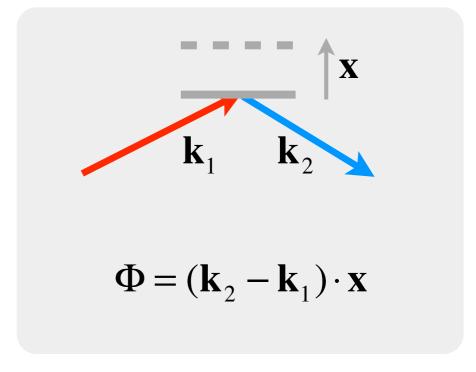
Why Atom Interferometers?

- Not because atoms have short de Broglie wavelength.
- But because atom clouds could be used as test masses for low frequency detection.
- Some recent proposals
 - R. Chiao et al.
 - S. Dimopoulos, P.W. Graham, J.M. Hogan, M.A. Kasevich and S. Rajendran, Phys. Rev. D 78, 122002, (2008).
- Paper on how atom interferometers respond to gravitational waves
 - A. Roura, D.R. Brill, B.L. Hu, C.W. Misner, and W.D. Phillips, Phys. Rev. D 73, 084018 (2006)

How does an atom interferometer work?

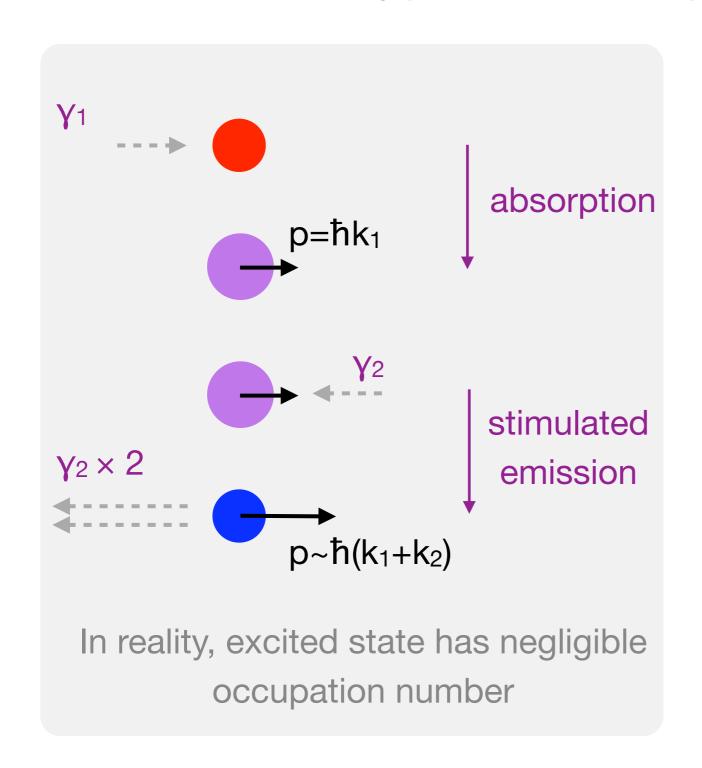


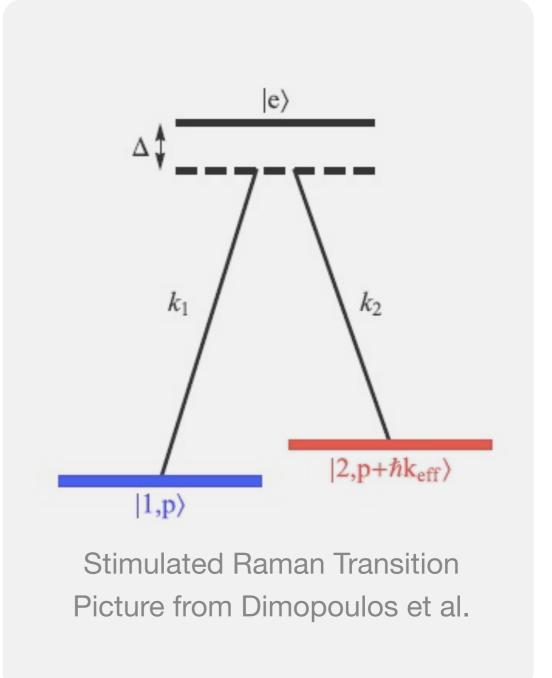
- Similar to optical interferometer
- Phaseshift depends on momentum transfer at mirror
- De Broglie wavelength $\lambda = h/(mv)$ could be small, de Broglie $k = mv/\hbar$ could be high, but it is δk that matters
- How do atom mirrors transfer momentum?



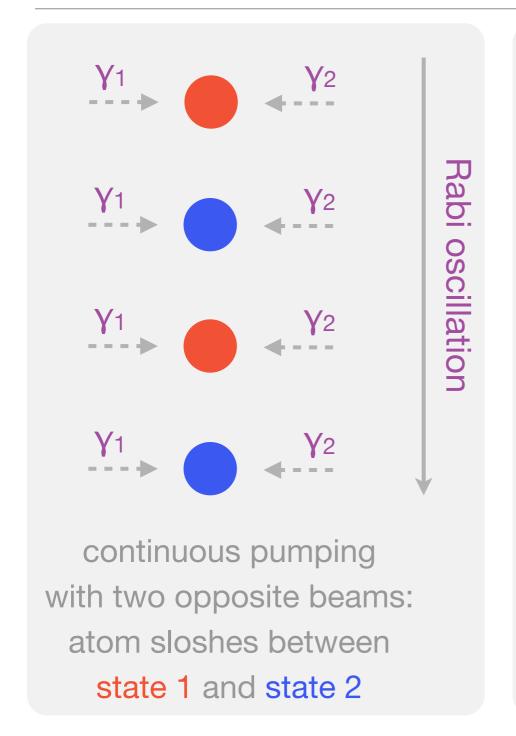
How does an atom mirror work?

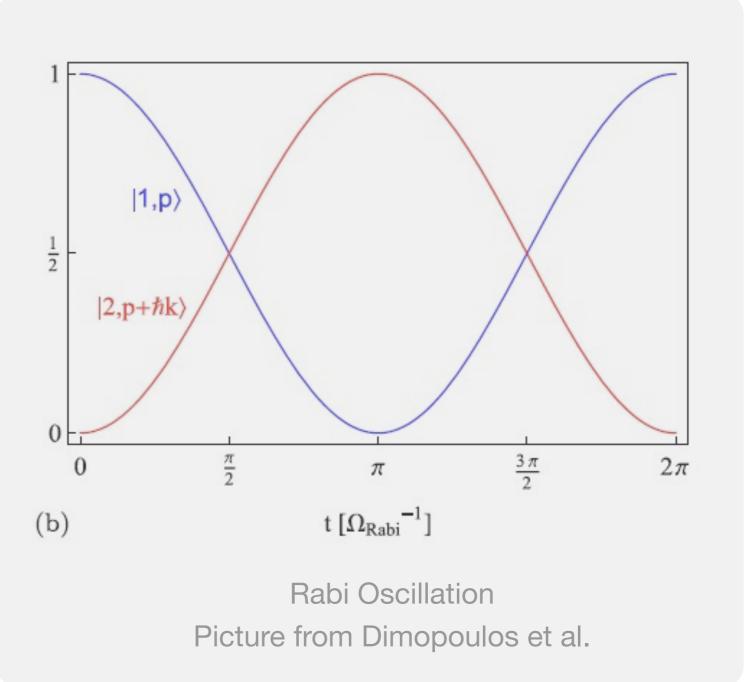
Atoms are deflected by photons, for example Stimulated Raman Transition





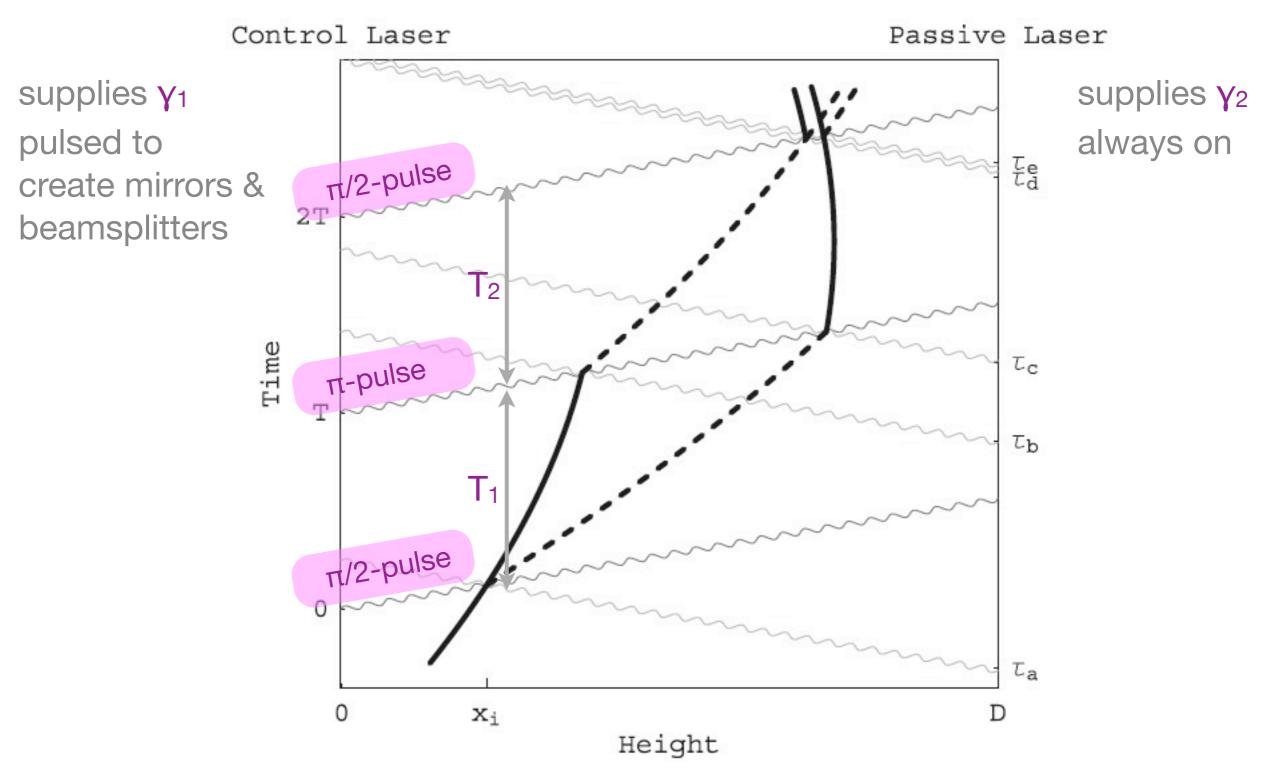
Mirror vs Beamsplitter: Phase of Rabi Oscillation





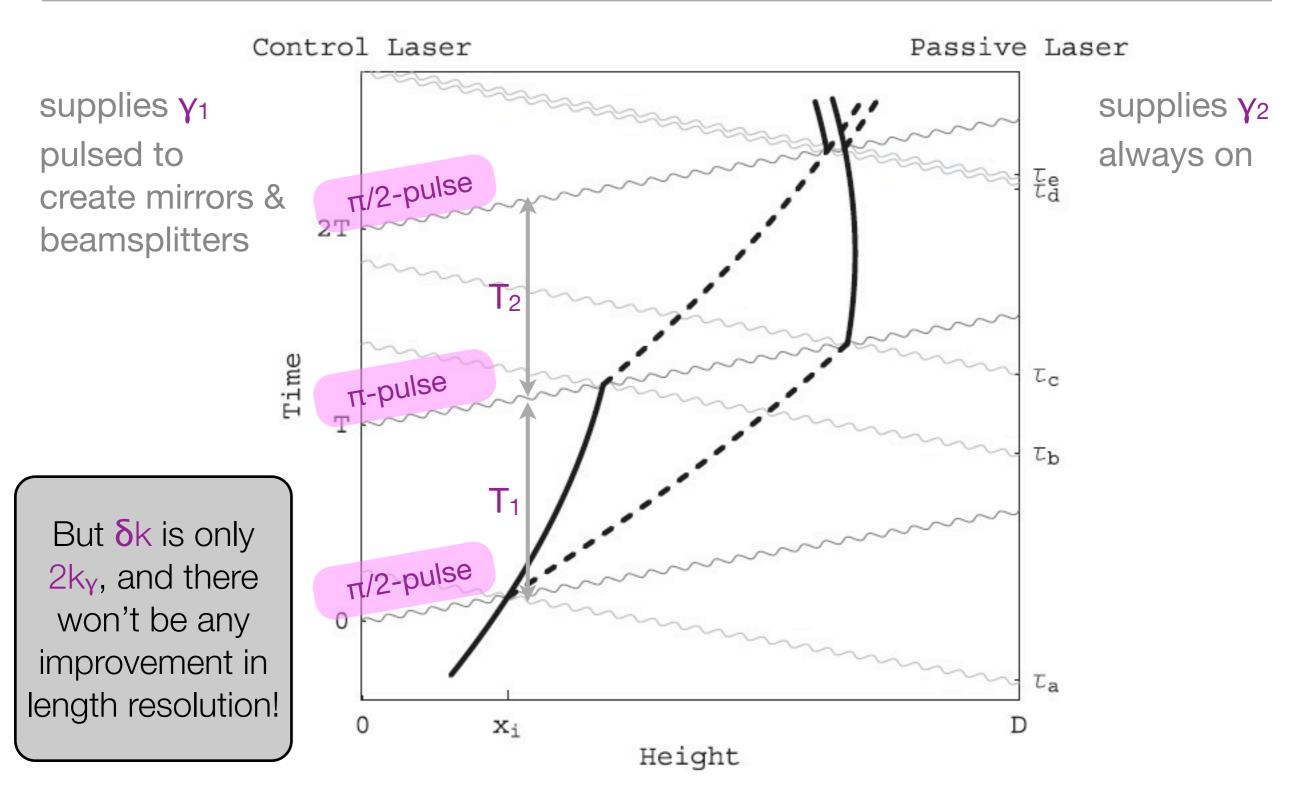
 Depending on phase of the Rabi oscillation, one can make a mirror (π-pulse) or beamsplitter (π/2-pulse)

An Atom Interferometer



Space-time diagram of a "linear" atom interferometer Picture from Dimopoulos et al.

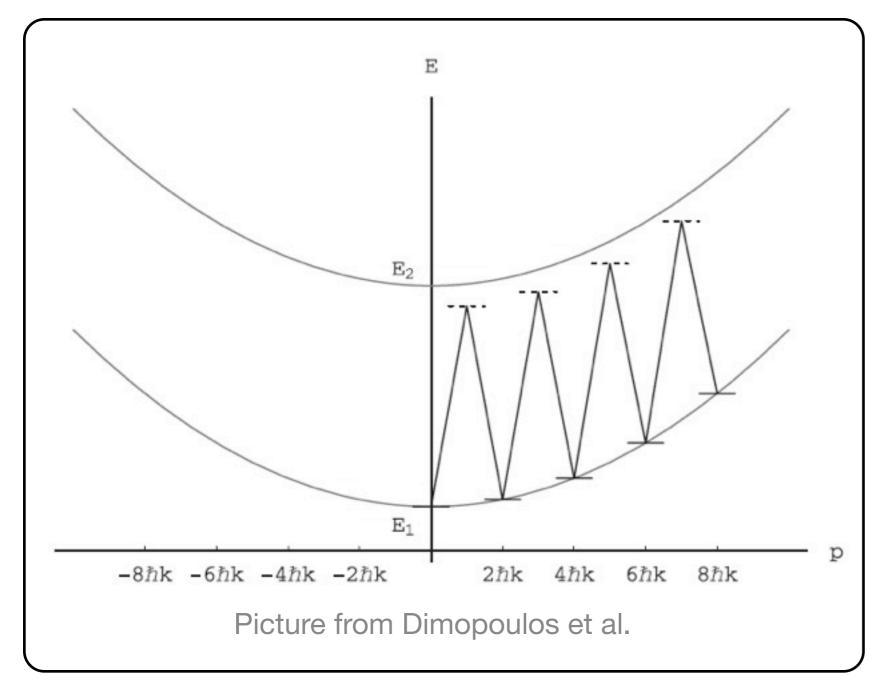
An Atom Interferometer



Space-time diagram of a "linear" atom interferometer Picture from Dimopoulos et al.

Large Momentum Transfer & Atom Squeezing

• It is possible to transfer more momenta, if the transition is Bragg, not Raman



Pulses can be iterated to give N times the momentum

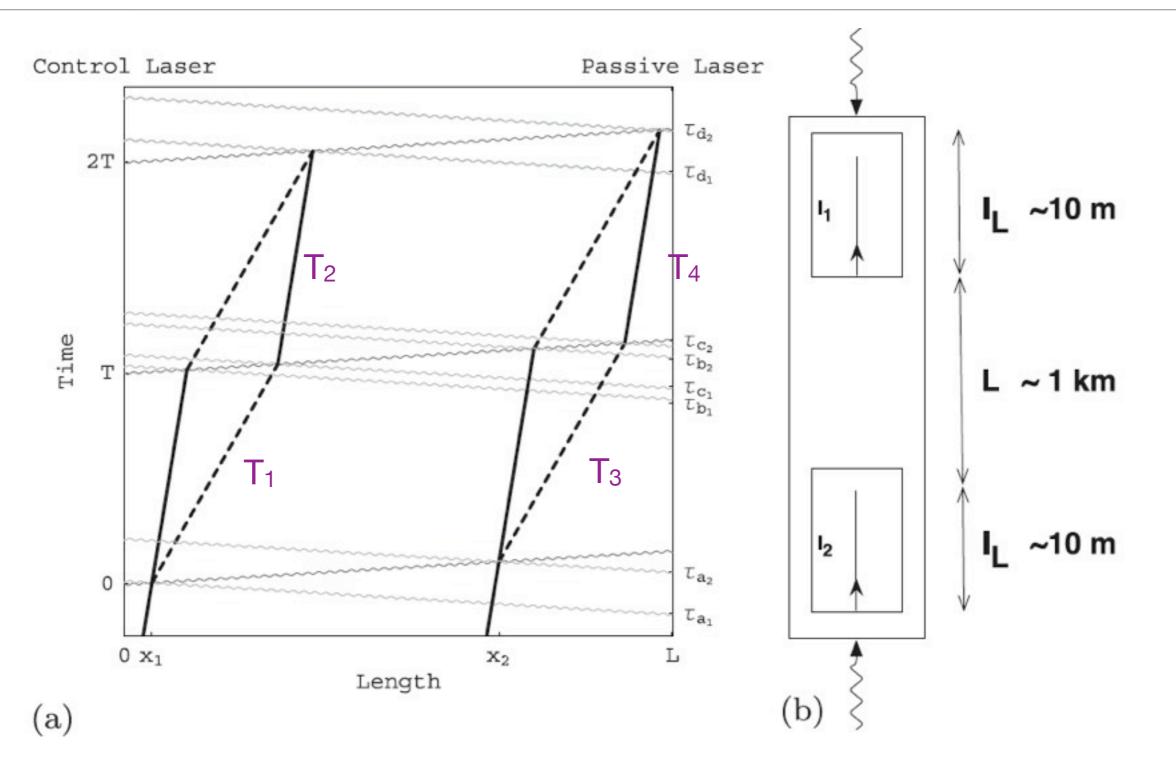
N ~ 100 - 1000 might be realizable

• Roughly equivalent to the use of a cavity with Finesse equal to *N*, which is not very high.

Why are we still interested in Atom Interferometry?

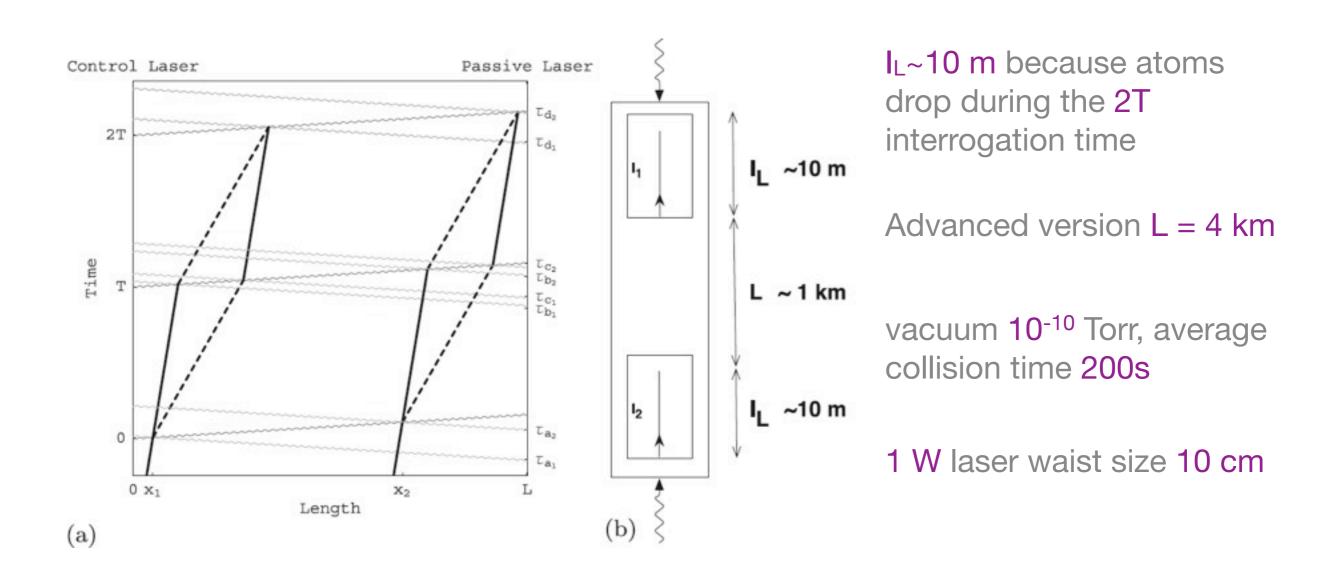
- Atoms might be usable as test masses for low-frequency detectors
 - they do not need to be suspended
 - they do not couple to vibration during flight
- Recent proposal
 - S. Dimopoulos, P.W. Graham, J.M. Hogan, M.A. Kasevich and S. Rajendran, Phys. Rev. D 78, 122002, (2008).
 - Two versions: terrestrial & space-based

Dimopoulos et al.'s Detector



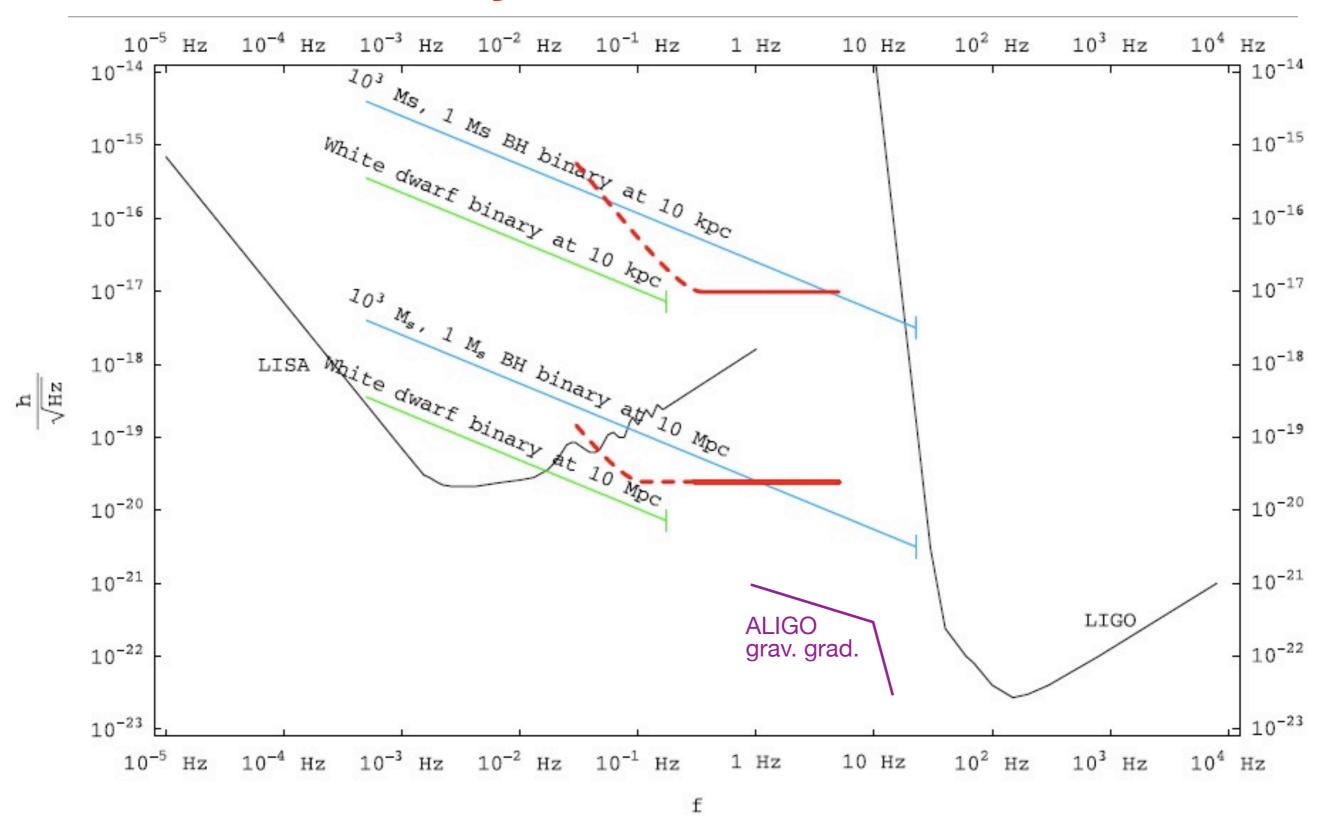
- Two co-linear atom interferometers sharing the same control and passive light
- Measures delays of laser pulses due to gravitational wave: (T₄-T₃)-(T₂-T₁)

Ground-based version

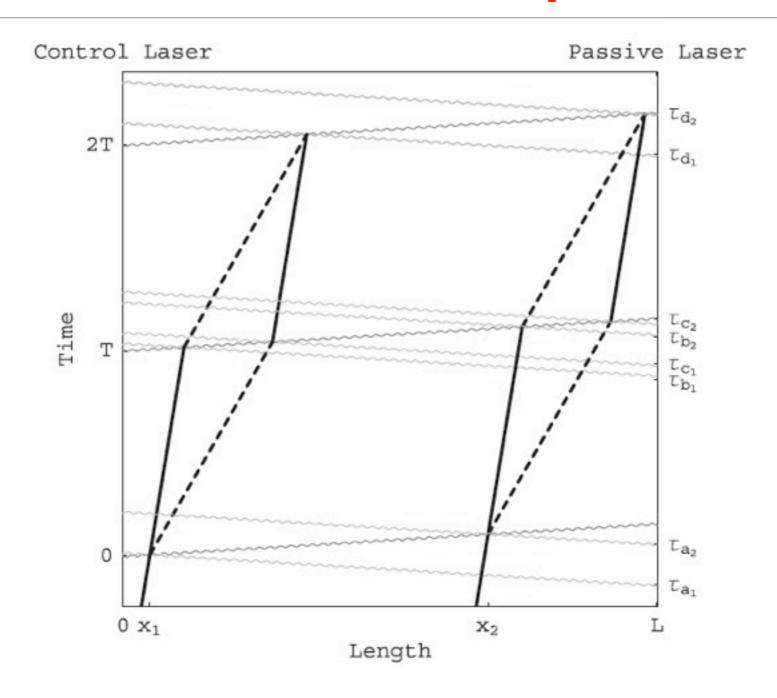


• $T \sim 1$ s, $k_{\text{eff}} \sim 1.6 \times 10^9$ (with $N \sim 1000$), $n \sim 10^8$, 5 clouds simultaneously, $h \sim 10^{-17}$ /rtHz at 1 Hz (advanced version is 10x better)

Sensitivity of Ground-Based Version

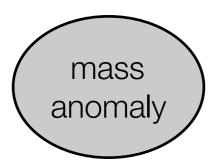


Noises: vibration & laser phase noise

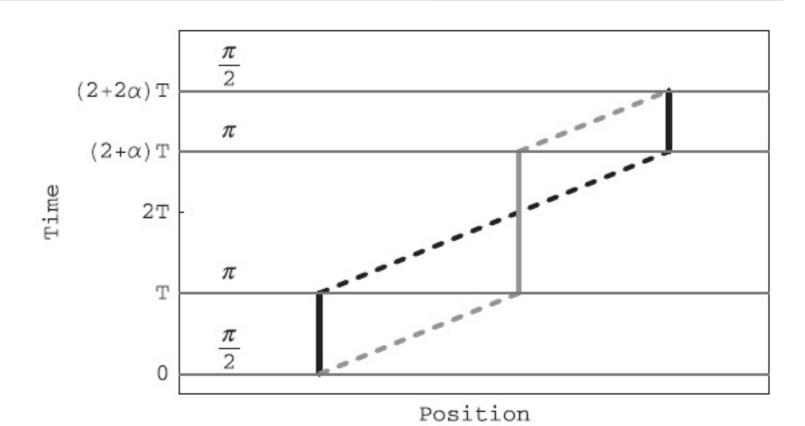


- Phase noise in control pulse unimportant (except for quantum part which seems small anyway)
- Vibration & Phase noise: phase noise of passive laser are subtracted between times of L/c (short). Requires $\delta x \sim 10^{-15} m$ (for 4 km version) or $10^{-14} m$ (for 1 km version)

Noises: Gravity Gradient



Mass anomaly may increase/ decrease gravity gradient



Different paths sense different **g**

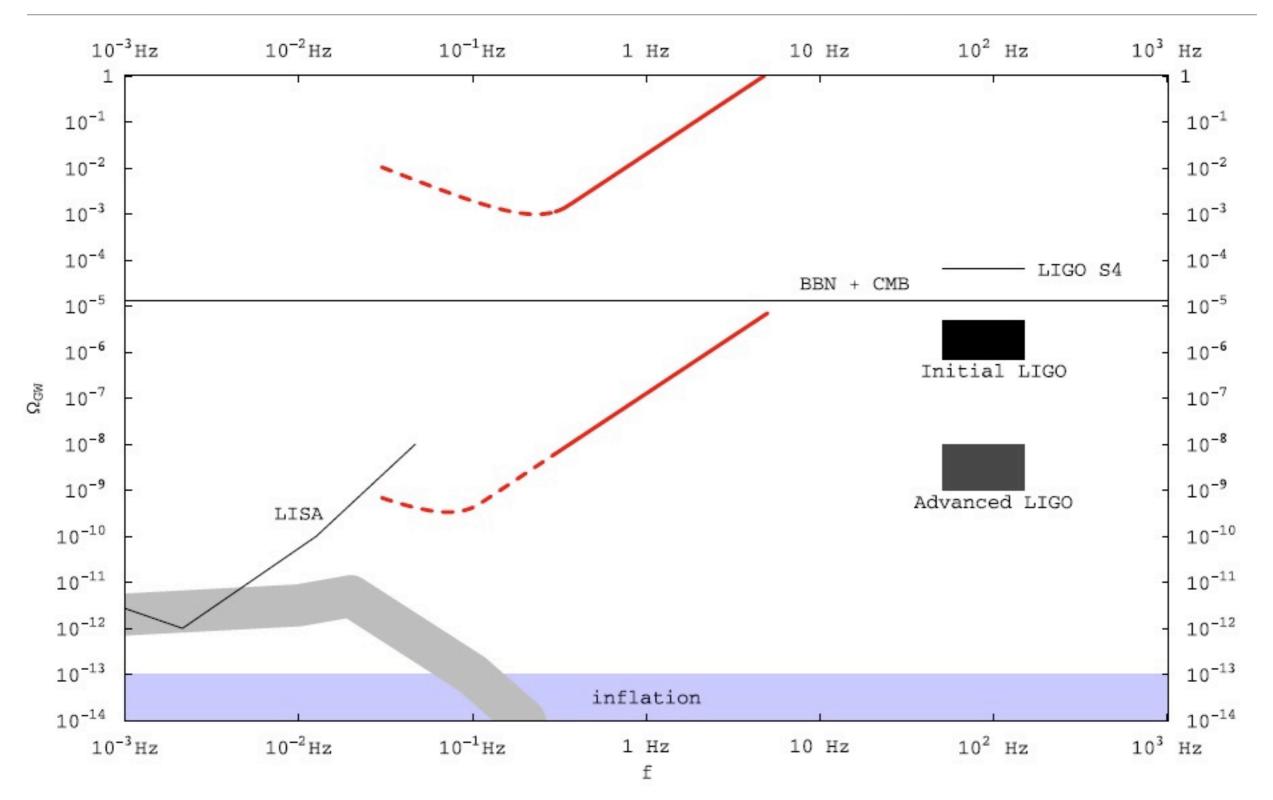
Double loop configuration, cancels g gradient

- Gravity Gradient: mostly due to gradient of g --- fluctuations in atoms' path cause difference in g. This sets requirement on the path of atoms, and therefore on the atom trap.
- Tricks can be used to avoid measuring them
 - Add compensation mass
 - Using fancy pulses
- Temporal fluctuations in gravity gradient seems unimportant.

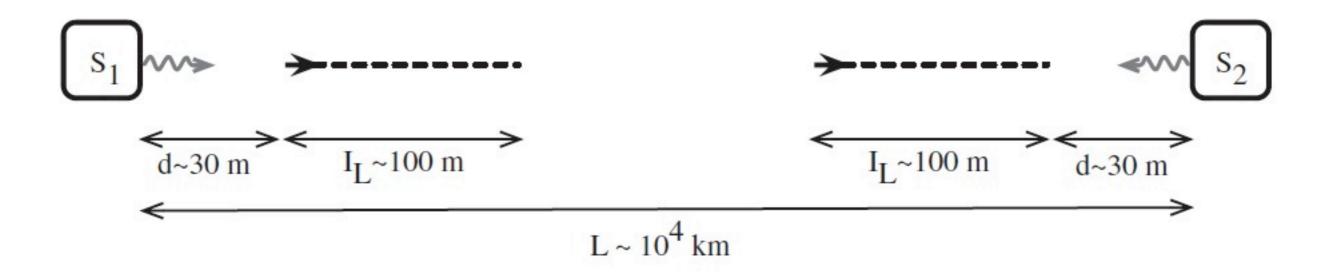
Rotation

- Coriolis force cause acceleration larger than GW $\delta a=2\omega_e\delta v$ with $\delta v\sim 10^{-8} m/s$
 - Can be cancelled if laser axis is not rotating together with earth. Requires accuracy of $\delta\omega$ <10⁻⁹(rad/s/rtHz)×(1 km/L)^{1/2}
- Laser jittering will cause the atoms to sense different ${\bf g}$, and this requires laser beam to have $\delta \omega < 10^{-9} (rad/s/rtHz)$ (for 1 km version)

Sensitivity to Stochastic Background



Space-Based Version



 $T \sim 100s$, below 0.01 Hz

d ~30 m to aiming at waves minimize gravity gradient caused by spacecraft

 $I_{I} \sim 100 \text{ m}$ is the maximum that spacecraft can provide Laser 1 W bias B field. For T ~ 100s, this limits k_{eff} to ~109. [V = $\hbar k/m$

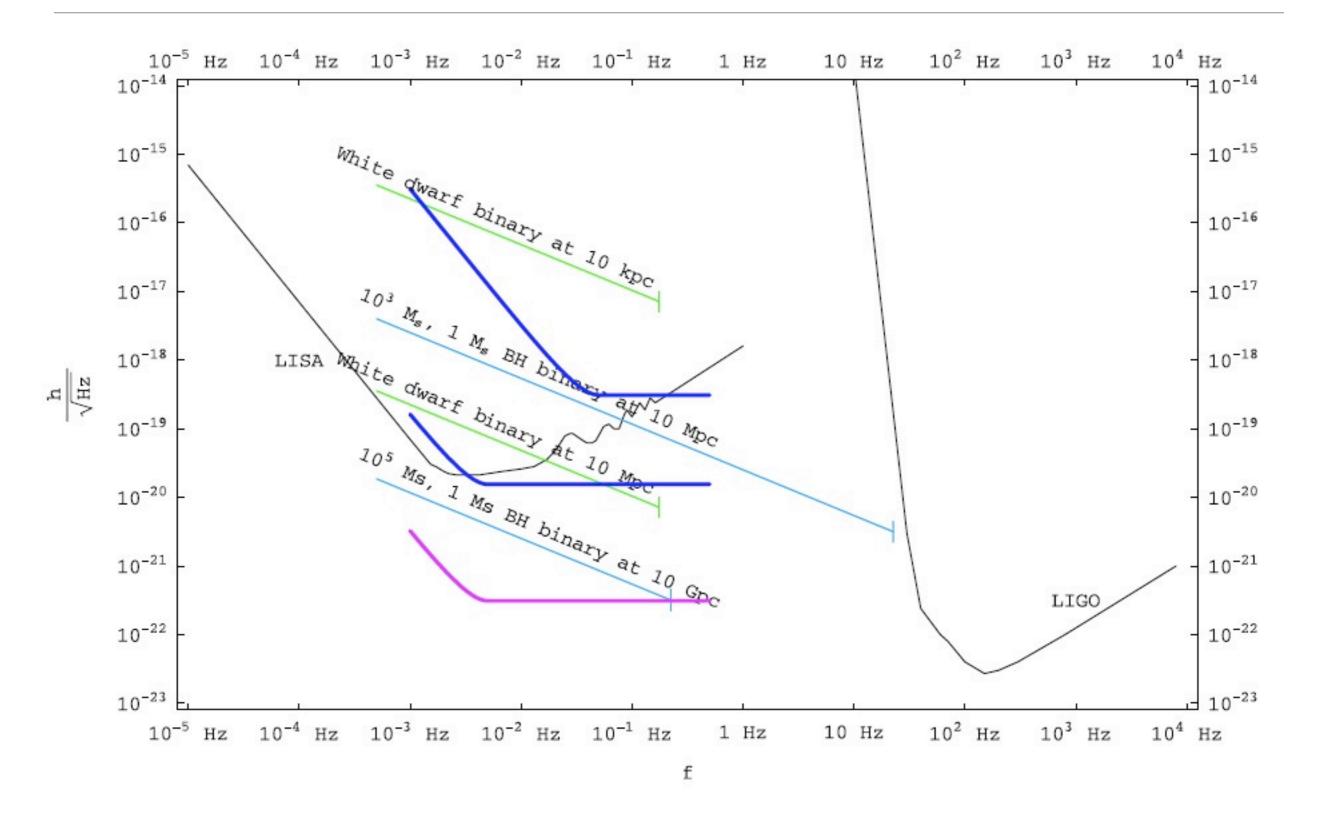
For 10³ km version, waistsize 0.5m

Must be shielded from sunlight, otherwise will only work when it's behind the earth, when photon-atom scattering timescale is longer than T

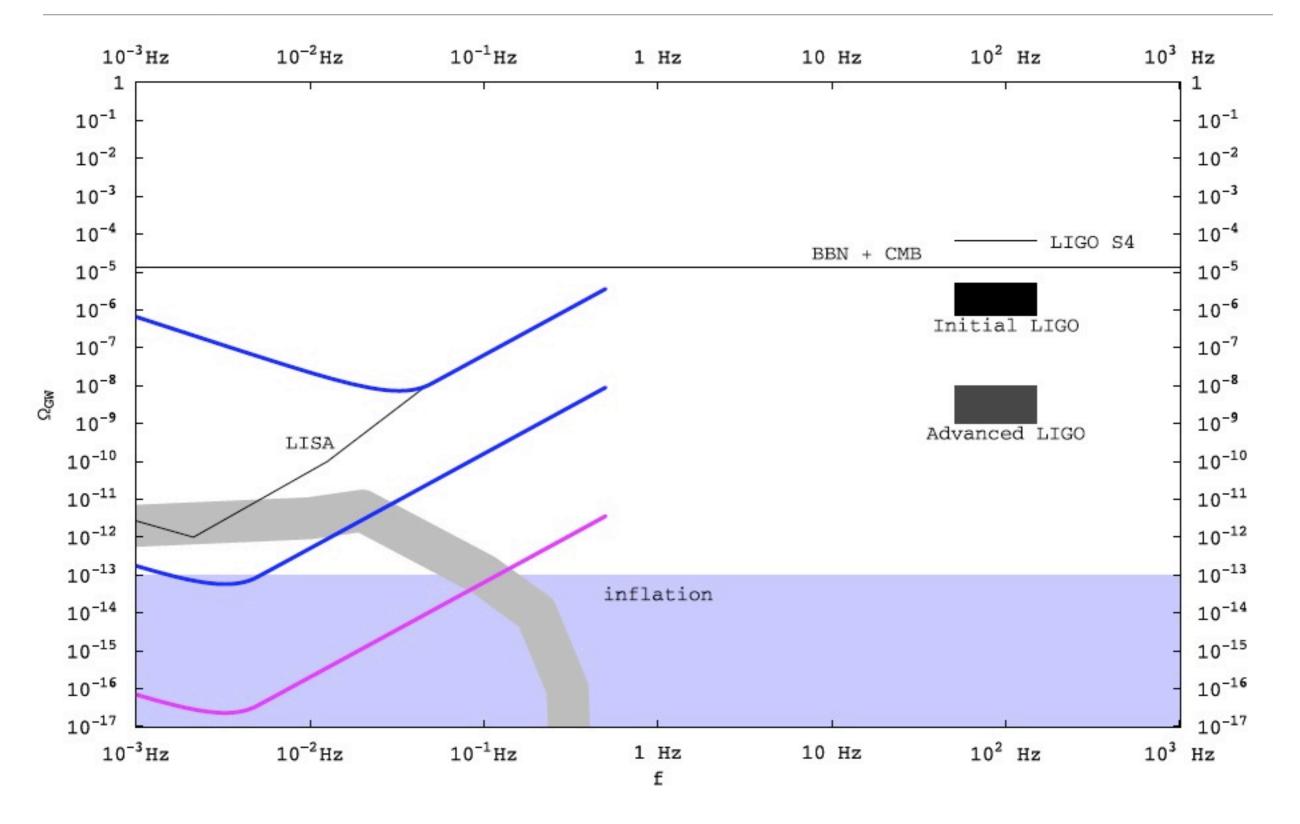
Effect from solar wind considered unimportant

Interplanetary magnetic field can be dealt with

Sensitivity of Space-Based Detector



Sensitivity to Stochastic Background



Comparison with LISA requirements

TABLE II. A comparison between specifications for a three-satellite AGIS configuration that could potentially allow comparable sensitivity to LISA, and the LISA requirements. There are many caveats and details that cannot be captured in a table and are discussed in Secs. VB and VC and in the LISA papers (see e.g. [42,47]).

I I I I I I I I I I I I I I I I I I I		
Attribute	AGIS	LISA
Baseline Satellite control (at ~10 ⁻² Hz)	10 ³ km	$5 \times 10^6 \text{ km}$
Laser frequency control (at $\sim 10^{-2} \text{ Hz}$)	$10^4 \frac{\text{nm}}{\sqrt{\text{Hz}}}$ $10^4 \frac{\text{Hz}}{\sqrt{\text{Hz}}}$	$ \begin{array}{c} 1 \frac{\text{nm}}{\sqrt{\text{Hz}}} \\ 1 \frac{\text{Hz}}{\sqrt{\text{Hz}}} \\ 1 \frac{\text{nrad}}{\sqrt{\text{Hz}}} \end{array} $
Rotational control (at $\sim 10^{-2}$ Hz) Electromagnetic forces	10 ⁻² nrad / Hz Atoms neutral, EM forces naturally small, predictable response to measured EM field	1 nrad √Hz Cosmic ray charging of proof mass
Collisions with background gas	Delete atoms, not a noise source	Cause acceleration noise
	Photon shot noise does not seem to have been taken into account	Limited by photon shot noise

Discussions?

